Truncated post-Newtonian neutron star model

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As a preliminary step towards simulating the binary neutron star coalescing problem, we test a post-Newtonian approach by constructing a single neutron star model. We expand the Tolman-Oppenheimer-Volkov equation of hydrostatic equilibrium by the power of c^{-2}, where c is the speed of light, and truncate at various orders. We solve the system using the polytropic equation of state with the index $\Gamma=5/3$, 2, and 3, and show how this approximation converges together with mass-radius relations. Next, we solve the Hamiltonian constraint equation with these density profiles as trial functions, and examine the differences in the final metric. We conclude that the second ''post-Newtonian'' approximation is close enough to describe a general relativistic single star. The result of this Brief Report will be useful for further binary studies.
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I. INTRODUCTION

Several Earth-based interferometers designed to detect gravitational waves have been recently constructed. Detectors such as the Laser Interferometric Gravitational Wave Observatory (LIGO), VIRGO, GEO, and TAMA are expected to begin operating within a few years (see, e.g., [1]). In order to extract gravitational waveforms from noisy data and to discuss physical parameters, it is essential to predict waveforms in advance by both analytical and numerical approaches.

Binary neutron star systems are one of the most plausible sources of gravitational waves. They emit energy through gravitational radiation, shrink their inspiral orbits gradually, and finally merge with strong emission of gravitational waves. The system is described by the post-Newtonian (PN) approximation (see, e.g., [2]) in the last several minutes before they merge, while in the last phase of coalescence of stars we need to solve the Einstein equations which are available only through numerical integration.

After the pioneering numerical works by Oohara and Nakamura in Newtonian gravity with a radiation reaction correction [3], several groups started developing numerical codes to solve this problem in a more realistic way. Such hydrodynamical simulations are categorized as in the Newtonian scheme (with or without a radiation reaction term) [4–11], post-Newtonian (PN) approximation [12], and fully general relativistic (GR) level [13–15]. However, we do not have a method to construct physically satisfactory initial data for inspiral binaries in general relativity. Most of the numerical tests start their simulations under assumptions of a certain quasiequilibrium and conformal flatness of spacetime, with a particular choice of vorticity of fluid (e.g., [16] and references therein).

One way to prepare initial data might be by patching the PN scheme to the general relativistic one [17]. In this Brief Report, we construct a simple model and examine how this effort is justified. We solve the Tolman-Oppenheimer-Volkov (TOV) equation of hydrostatic equilibrium of a single neutron star, which is truncated at the various PN levels. We compare the mass and radius of a star as a function of central density using the polytropic equation of state. We also solve the Hamiltonian constraint equation of the Einstein equations by substituting these density profiles as trial functions, and discuss the differences in the metric.

This study is an extended one from earlier works [18–21] using the first PN approximation. We intend to make a bridge between the Newtonian and general relativistic solutions of a neutron star model, both of which were first shown numerically by Tooper [22].

In the actual calculations, we used geometrical units of $c=G=M_{\odot}=1$, where c, G, and $M_{\odot}$ are the speed of light, Newton's gravitational constant, and the solar mass, respectively. However, c and G will appear in the text where they help understanding.

II. TRUNCATED TOV NEUTRON STARS

In general relativity, we have the TOV equation for solving a hydrostatic equilibrium star in spherically symmetric spacetime. We start from the metric

$$ds^2=-e^{2\Phi(r)}dt^2+e^{2\Lambda(r)}dr^2+r^2(d\theta^2+\sin^2\theta d\varphi^2), \tag{2.1}$$

where $e^{2\Lambda(r)}=[1-2Gm(r)/c^2r]^{-1}$. Then the TOV equations are written as

$$\frac{dm}{dr}=4\pi r^2\rho_t, \tag{2.2}$$

$$\frac{dp}{dr}=-\frac{Gm\rho_t}{r^2}\left(1+\frac{p}{\rho_t c^2}\right)\left(1+\frac{4\pi p r^3}{mc^2}\right)\left(1-\frac{2Gm}{rc^2}\right)^{-1}, \tag{2.3}$$

$$\frac{d\Phi}{dr}=-\frac{1}{\rho_t}\frac{dp}{dr}\left(1+\frac{p}{\rho_t c^2}\right)^{-1}, \tag{2.4}$$

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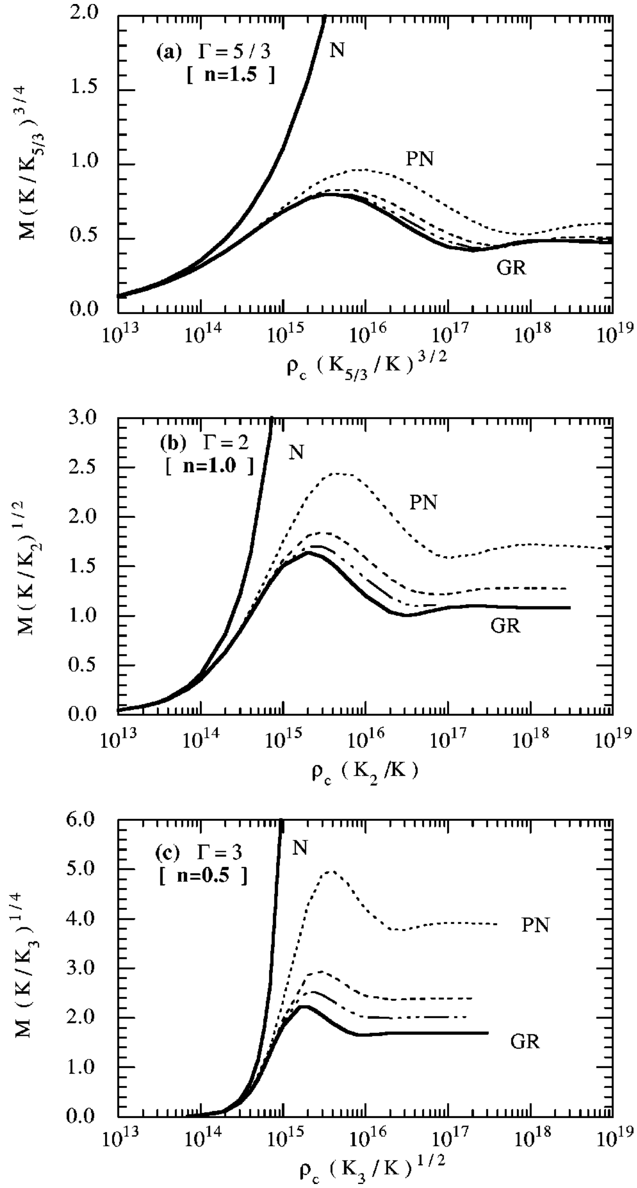

FIG. 1. Total mass as the function of its central density for truncated neutron star model. (a), (b), and (c) are for different equations of state with $\Gamma=5/3$, 2, and 3, respectively. Mass is in units of the solar mass and the central density is in [g/cm^3]. The gray solid line is of Newtonian solutions; the solid line is of general relativistic solutions. The dotted line, dashed line, and three-dotted line are of first, second, and third post-Newtonian approximated solutions, respectively.

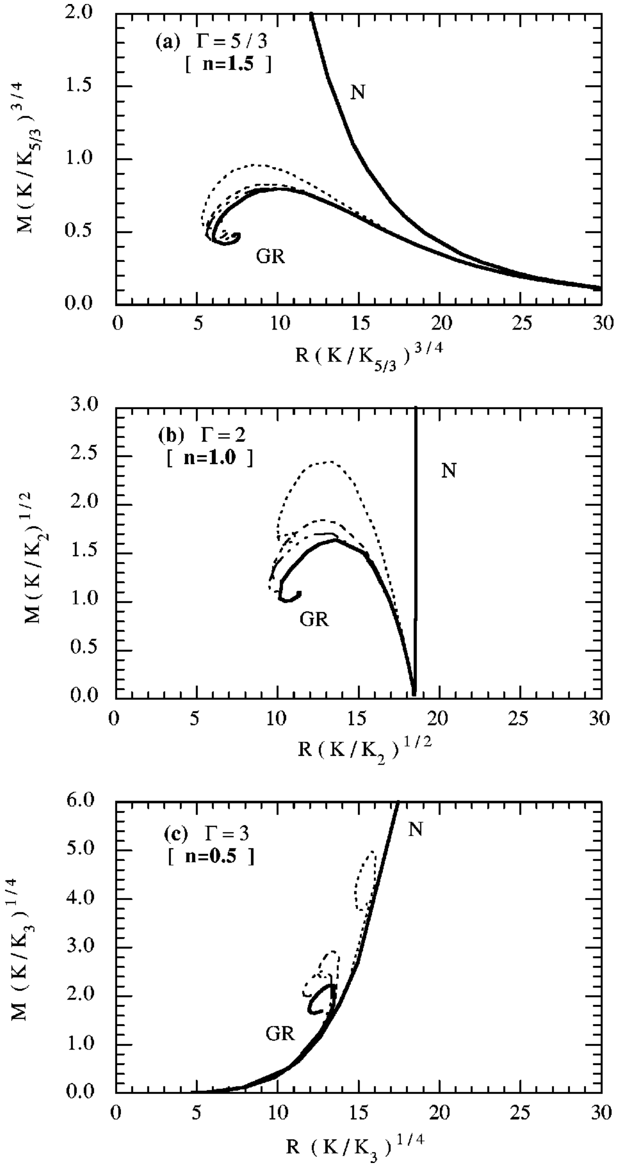

FIG. 2. Mass and radius relations for truncated neutron star models. Mass is in units of the solar mass and radius is in [km]. The lines are the same as in Fig. 1.

together with the specified equation of state, for which we use the polytropic equation of state

$$p=K\rho^{\Gamma}=K\rho^{1+1/n}, \tag{2.5}$$

where p and ρ are the pressure and energy density, respectively, and ρ_t is the total mass density:

$$\rho_t=\rho+\frac{p}{(\Gamma-1)c^2}. \tag{2.6}$$

Obviously, the set of equations recovers the Newtonian limit for $c^2\to\infty$.

The idea of this Brief Report is to expand the product of the parentheses in Eqs. (2.3) and (2.4) and truncate them at the order of $1/c^{2i}$. The ith truncation, then, gives the so-called ith PN approximation. (The case of $i=1$ is briefly mentioned in [23].) That is, we write Eqs. (2.3) and (2.4) schematically:

$$\begin{aligned}\frac{dp}{dr}&=-\frac{Gm\rho_t}{r^2}(1+A)(1+B)(1-C)^{-1}\\&=-\frac{Gm\rho_t}{r^2}(1+A+B+C\\&\quad+AB+AC+BC+C^2+\cdots),\end{aligned} \tag{2.7}$$

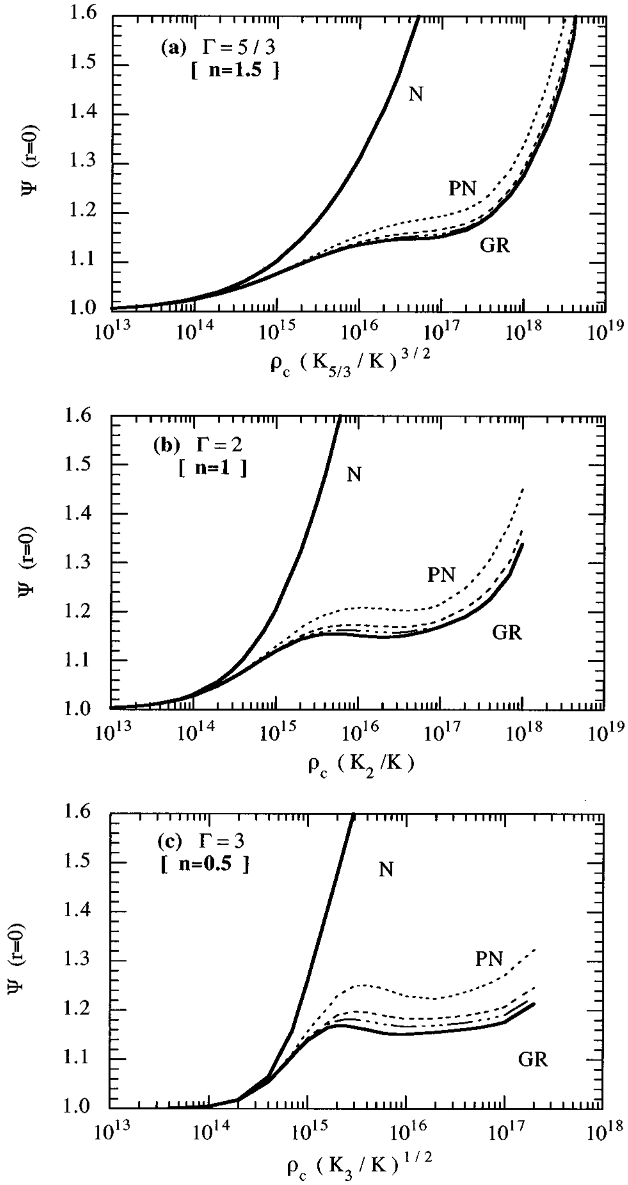

FIG. 3. The conformal factor ψ at the origin is displayed as a function of central density, of which we used a trial configuration for solving Hamiltonian constraint equation. The central density is in units of [g/cm^3]. Each line indicates the trial profile as input, using the same notation as Fig. 1.

$$\frac{d\Phi}{dr} = -\frac{1}{\rho_t}\frac{dp}{dr}(1+A)^{-1}$$

$$= -\frac{1}{\rho_t}\frac{dp}{dr}(1-A+A^2-A^3+\cdots). \tag{2.8}$$

If we use these equations with terms on the right-hand side (RHS) of up to two products of A,B,C (such as AB or A^2), then we say the system is in the second PN approximation.

We apply $\Gamma=5/3,\ 2$, and 3 for the equation of state ($n=1.5,\ 1$, and 0.5 in the polytropic index, respectively) and compare the solutions of Newtonian, GR, and up to the third PN approximation.

The radius of the star, R, is measured at the point $r_\star$ where density ρ_t drops low enough [$O(10^{-10})$ in geometrical units], and given by the proper length,

$$R = \int_0^{r_\star}\left(1-\frac{2Gm(r)}{c^2 r}\right)^{-1/2} dr, \tag{2.9}$$

with appropriate truncation in the integrand. We express the mass of the star, M, by $M=m(r_\star)$.

We use the fifth-order Runge-Kutta method (Fehlberg method) to integrate the equations. In order to check that this approach is right, we also worked the TOV equations in the harmonic gauge and confirmed that we get identical physical quantities in the results.

In Fig. 1, we show the total mass M as a function of the central density ρ_c for the different Γ's and PN levels. Mass is in units of $M_\odot$ and central density is in [g/cm^3], and both are rescalable with the constant K in the equation of state. Here we use K in the calculations as $K_{5/3}=4.35$ (for $\Gamma=5/3$), $K_2=10^2$ (for $\Gamma=2$), and $K_3=10^5$ (for $\Gamma=3$) in geometrical units, where $K_{5/3}$ is the number for the pure neutron equation of state [24].

We see clearly the convergence of this PN approximation in all the Γ's. However, if the equation of state is stiff, then the high density configuration differs from that of GR even at the higher PN approximation.

From the first PN approximation, we see the existence of the maximum mass. The central density which gives this maximum becomes larger in the weak gravity approximation.

In Fig. 2, we show the mass-radius relations. In the Newtonian limit, the asymptotic behaviors of M near $M=0$ are as $M\propto R^{-3}$ (for $\Gamma=5/3$), $M\propto R^0$ (for $\Gamma=2$), and $M\propto R^5$ (for $\Gamma=3$). These represent the softness (for $\Gamma=5/3$) and stiffness (for $\Gamma=3$) of the equation of state. We see that all the lines in Fig. 2 coincide with this Newtonian limit in the lower mass limit. The figure also shows us that the first PN solution has the same feature as GR.

We also checked the causality constraint $dp/d\rho\leqslant 1$ (see, e.g., [25]) in all of the models, and confirmed that the constraint is always valid.

III. METRIC OUTPUT VIA THE HAMILTONIAN CONSTRAINT

We next solve the Hamiltonian constraint equation in GR with the trial density profiles obtained above. Our aim is to compare the difference of the output metric and to examine a matching scheme of PN data to the general relativistic one.

We use the O'Murchadha-York conformal approach [26] to solve the Hamiltonian constraint. Defining the conformal factor ψ and setting $\gamma_{ij}=\psi^4\hat{\gamma}_{ij}$, the constraint becomes

$$8\,{}^{(3)}\hat{\Delta}\psi = {}^{(3)}\hat{R}\psi - 16\pi G\hat{\rho}\psi^{-3}, \tag{3.1}$$

where ${}^{(3)}\hat{\Delta}$ and ${}^{(3)}\hat{R}$ are the three-dimensional Laplacian and Ricci scalar curvature, respectively, defined by $\hat{\gamma}_{ij}$. Here we assumed $K_{ij}=\hat{K}_{ij}=0$.

We choose our trial metric $\hat{\gamma}_{ij}$ as conformally flat, and solve Eq. (3.1) with a trial density configuration of $\hat{\rho}=\rho_t$. We use the incomplete Cholesky conjugate gradient (ICCG) method [27] with the Robin boundary condition $\psi=1+C/r$, where C is a constant, for solving Eq. (3.1).

In Fig. 3, we show the conformal factor ψ at the origin as a function of the central density of the trial configuration. The three-metric at the center will be given by $\gamma_{ij}=\psi^4\delta_{ij}$. We see that using the Newtonian configuration as input gives us quite different solutions from the expected ones of GR, while all PN trials give similar solutions with GR. Independently of Γ, we can say that the second PN approximation provides closer values for the output metric to those of GR.

IV. DISCUSSION

In order to justify the recent post-Newtonian approaches to the binary neutron star problem, we constructed a simple model. By solving the hydrostatic equilibrium equation of a star at the *i*th PN approximation, we showed the convergence of this approach, the mass and radius relations, and resultant metric output via the Hamiltonian constraint equation.

We conclude that the second PN approximation provides quite similar density profiles to those of GR, independent of equations of states. If we use second PN density configurations as trial functions, we get closer metric solutions to those from GR through the Hamiltonian constraint. Although this study is restricted to a hydrostatic single star model, we think that the figures shown here are convenient templates for further numerical studies.

As shown in [17], the discontinuous matching surface of PN and GR in the vacuum region will be smoothed out in fully relativistic evolution in a particular slicing condition. Therefore we expect that higher PN initial data will smoothly evolve in the fully relativistic simulations, although there are many unknown factors as to whether such initial data are numerically satisfactory or not. We are now applying this approach to construct a binary model including their velocity corrections together with fully general relativistic hydrodynamical evolutions. This effort will be reported elsewhere.

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